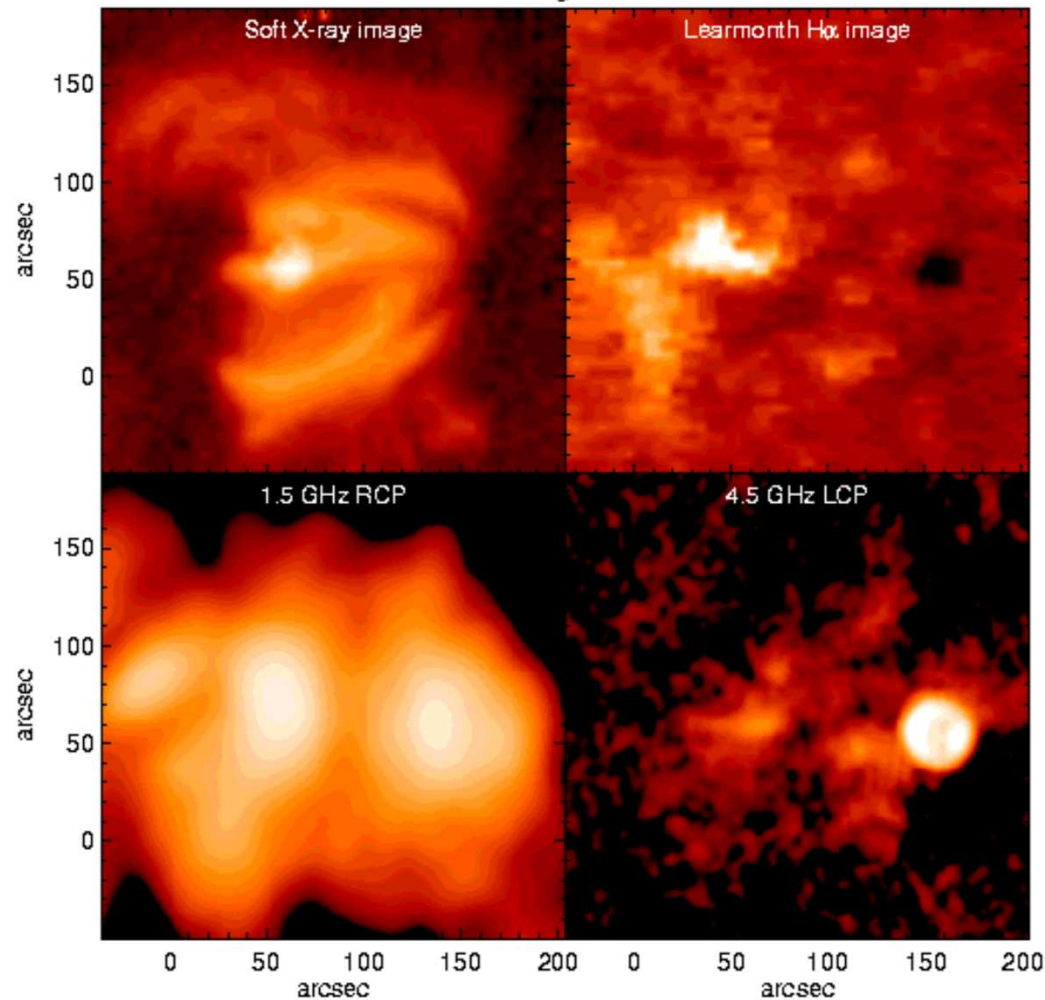


For typical coronal conditions, *X* mode is optically thick in $s=2$ & 3 over a broad range of angles θ . The *O* mode is optically thick over most of $s=2$, and may be at least marginally optically thick over a small portion of $s=3$ if θ is large. Harmonics greater than $s=4$ do not have any significant optical depth in the quiet solar corona.

Fig. 5.1 in Gary & Keller (2004)

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Gyroresonance emission is always dominant at frequencies above a few GHz in strong-field regions above sunspots, while free-free emission is generally dominant in the weak-field regions in plages of ARs, almost always at frequencies of $\nu \gtrsim 2$ GHz. Measuring the circular polarization at $\nu \gtrsim 1$ GHz provides information on the magnetic field strengths for both emission mechanisms.

Fig. 5.27 in Markus J. Aschwanden (2005)